

CONSTRAINTS ON THE CLUSTER ENVIRONMENTS AND HOT SPOT MAGNETIC FIELD STRENGTHS OF THE RADIO SOURCES 3C 254 AND 3C 280

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ABSTRACT

We present new *Chandra Observatory* observations together with archival *Hubble Space Telescope (HST)* and radio observations of 3C 254, a radio quasar at $z = 0.734$, and 3C 280, a radio galaxy at $z = 0.996$. We report the detection of X-ray and possible *HST* optical counterparts to the radio hot spots in 3C 280 and the detection of an X-ray counterpart to the western radio hot spot in 3C 254. We present constraints on the presence of X-ray clusters and on the magnetic field strengths in and around the radio hot spots for both targets. Both sources were thought to be in clusters of galaxies based on reports of significant extended emission in *ROSAT* PSPC and HRI images. The exquisite spatial resolution of the *Chandra Observatory* allows us to demonstrate that these sources are not in hot, massive clusters. The extended emission seen in *ROSAT* observations is resolved by *Chandra* into point sources and is likely to be X-ray emission associated with the radio hot spots of these sources, with possible additional contributions from unrelated point sources. The intergalactic medium around these sources could be dense, but it is demonstrably not dense and hot. We conclude that radio sources are not reliable signposts of massive clusters at moderately high redshifts. We present measurements of the X-ray and optical fluxes of source features and discuss what physical processes may give rise to them. X-ray synchrotron emission could explain the radio, optical, and X-ray hot spot fluxes in 3C 280; this would require continuous acceleration of electrons to high Lorentz factors since the synchrotron lifetime of relativistic electrons that could produce the X-ray emission would be of the order of a human lifetime. Synchrotron self-Compton emission with or without inverse Compton emission due to scattering of the cosmic microwave background radiation can also explain the X-ray emission from the hot spots, although this most likely would require that some other physical process be invoked to explain the optical emission seen in 3C 280. High spatial resolution radio data with broad frequency coverage of the radio hot spot regions are needed to determine which physical process is responsible for the detected X-ray emission and to provide much tighter constraints on the magnetic field strength of the hot spot plasma. We summarize our current constraints on the magnetic field strengths in and around the hot spots of 3C 254 and 3C 280.

Subject headings: cosmology: observations — galaxies: active — galaxies: clusters: general — galaxies: individual (3C 280) — galaxies: magnetic fields — quasars: individual (3C 254)

On-line material: color figures

1. INTRODUCTION

Since radio sources are visible from across the universe, understanding the connections between their observed properties and their physical conditions allows us to probe the environment and physics of galaxies billions of years ago. Radio sources are thought to inhabit regions of high galaxy density because of an excess of companions (see, e.g., Pascarella et al. 1996; Le Fèvre et al. 1996) and large rotation measures that could arise from hot, magnetized cluster gas (Carilli et al. 1997). Astronomers have speculated that the most distant radio sources evolve into massive elliptical galaxies (see, e.g., Pentericci et al. 1999). De Vries et al. (2000) show that the galaxy hosts are likely to be giant elliptical galaxies. Best et al. (1997a, 1997b) showed that magnitudes, colors, and profiles of 3CR radio galaxies at $z \sim 1$ are consistent with those of an $M \sim 5 \times 10^{11} M_{\odot} \text{ yr}^{-1}$ elliptical galaxy with a formation redshift of 3–5. The closest example of such a system is Cygnus A. For reference, Cyg A is a powerful radio galaxy inhabiting and interacting with the

intracluster medium (ICM) of a relatively rich cluster of galaxies with an X-ray bolometric luminosity of $L_X \sim 6.2 \times 10^{44} \text{ ergs s}^{-1}$ (converted from a 2–10 keV luminosity of $L_X \sim 3.5 \times 10^{44} \text{ ergs s}^{-1}$) and X-ray temperature $T_X \sim 7.7 \text{ keV}$ (Smith et al. 2002).

In this work, we observed two powerful FR IIb radio sources. An FR IIb source is defined as an edge-brightened, classical double radio source with a quite regular (cigar-like) radio bridge region, and it is also classified as a type 1 FR II source by Leahy & Williams (1984). FR IIb sources, described in some detail by Daly (2002), typically have radio powers at least a factor of 10 above the classical FR I–FR II break and are found almost exclusively at relatively high redshift (with the well-known exception, Cyg A). For FR IIb sources, it is thought that the ram pressure confinement of the radio lobe might be used to estimate the ambient gas density using radio data alone, as was done by Carilli et al. (1991) for Cyg A, Perley & Taylor (1991) for 3C 295, Wellman, Daly, & Wan (1997, hereafter WDW97) for a sample of 22 FR IIb sources, and Guerra, Daly, & Wan (2000) for six FR IIb sources. A comparison of radio-predicted ambient gas densities with X-ray-determined ambient gas densities suggests that the magnetic field strength in the radio bridges of 3C 295 and Cyg A are about $\frac{1}{4}$ of the minimum-energy value (Perley & Taylor 1991; Carilli et al. 1991).

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If the ram pressure method of determining ambient gas densities were confirmed to be reliable by an X-ray assessment of the ambient gas densities, then we might be able to use the radio data to ascertain the density of the ICM at higher redshifts where the X-ray emission is nearly impossible to detect with current technology. Such a method would be complementary to Sunyaev-Zeldovich (SZ) techniques since, in general, SZ imaging is more difficult near powerful radio sources. The properties of the ICM at high redshifts, insofar as they reveal the virial mass of the cluster, provide important constraints on models of evolution of clusters and of large-scale structure, strongly constraining Ω_{matter} (see, e.g., Peebles, Daly, & Juskiewicz 1989; Donahue et al. 1998, 1999; Donahue & Voit 1999; Henry 2000, 1997; Eke et al. 1998; Borgani et al. 2001).

In order to test this method of assessing densities, we observed two very powerful, extended, classical double radio FR IIb sources, 3C 254 and 3C 280. Both of these sources had strong indications of extent in X-ray images from studies using *ROSAT* (Hardcastle & Worrall 1999; Crawford et al. 1999). The excellent spatial resolution of the *Chandra X-Ray Observatory* allows us to distinguish between the compact emission from a central point source and hot spots from extended emission from a thermal, extended gaseous halo gravitationally confined by the potential associated with a cluster of galaxies. Typical clusters of galaxies have ICM luminosities of $L_X \sim 10^{43}$ – 10^{45} ergs s^{-1} and gas temperatures of $T_X \sim 2$ – 15 keV. A hot ICM has been detected around at least some distant 3CR radio galaxies using *Chandra's* imaging detectors. 3C 294 at $z = 1.786$ was found to have diffuse emission of ~ 5 keV and a rest-frame 0.3–10 keV X-ray luminosity of $4.5 \times 10^{44} h_{50}^{-2}$ ergs s^{-1} (Fabian et al. 2001). 3C 295 ($z = 0.47$) has a 5 keV ICM, detected out to $1 h_{50}^{-1}$ Mpc with *Chandra* (Allen et al. 2001). 3C 220.1 ($z = 0.62$) is contained within a 5 keV ICM detected with *Chandra* with a rest-frame X-ray luminosity (0.7–12.0 keV) of 5.6×10^{44} ergs s^{-1} (Worrall et al. 2001). Despite the detections of clusters around other radio sources and the early assessments of extent in 3C 280 and 3C 254, in this work we find that neither of these sources is in a cluster of any significance ($L_X \gtrsim 3 \times 10^{43}$ ergs s^{-1}). We discuss our X-ray observations, our limits on extended X-ray flux, and what inferences may be made regarding the ICM surrounding these sources.

However, we also detect significant emission from X-ray counterparts of the radio hot spots in both sources. In this paper, we describe the two targets of our studies in § 2. We outline the details of the *Chandra* observations and analysis in § 3. We briefly discuss our reduction and analysis of *Hubble Space Telescope* (*HST*) data in § 4. The limits placed on the existence of the X-ray clusters and the density profiles of each system are contained in §§ 5.1 and 5.2, along with a review of the archival *ROSAT* data in § 5.3. In § 6.1, we present useful relations for computing the expectations for synchrotron self-Compton (SSC) processes in hot spots and apply them to our observations. In § 6.2 we discuss the contribution of synchrotron to the X-ray and possible optical emission; in § 6.3 we do the same for inverse Compton (IC) scattering of the microwave background. In all sections we place constraints on the magnetic field. In § 7, we summarize our conclusions regarding the physical environments of the radio sources, the physical processes causing the hot spots, and the magnetic fields in those hot spots.

For the purposes of comparing luminosities and length scales between this work and previous X-ray research, we assume $H_0 = 50 h_{50}$ km s^{-1} Mpc $^{-1}$ and $q_0 = 0.0$, implying an angular distance scale of $10.885 h_{50}^{-1}$ kpc arcsec $^{-1}$ for 3C 280 and $9.700 h_{50}^{-1}$ kpc arcsec $^{-1}$ for 3C 254. For reference, here are the angular distance scales and luminosity conversion factors for the currently most fashionable cosmology of $H_0 = 65$ km s^{-1} Mpc $^{-1}$, $\Omega_M = 0.3$, and $\Omega_\Lambda = 0.7$: the angular distance scales are $8.616 h_{50}^{-1}$ kpc arcsec $^{-1}$ for 3C 280 and $7.838 h_{50}^{-1}$ kpc arcsec $^{-1}$ for 3C 254. The luminosities quoted in this paper would decrease by a factor of the luminosity distance D_L^2 : 0.63 for 3C 280 and 0.65 for 3C 254.

2. INDIVIDUAL SOURCES

3C 254 is a radio-loud FR II quasar with $z = 0.734$ (Spinrad et al. 1985) at $11^{\text{h}}14^{\text{m}}38^{\text{s}}.5$, $+40^{\circ}37'20''$ (J2000.0). The radio source is very asymmetric, with its shortest hot spot–nuclear separation on the eastern side. Owen & Puschell (1984) showed that the eastern lobe is extended roughly perpendicular to the source axis. An overdensity of galaxies near the source suggests that the environment is a cluster or a group (Bremer 1997). This overdensity consists of seven objects within $10''$ ($97 h_{50}^{-1}$ kpc), with magnitudes within 3–5 mag of that of the central galaxy. The magnitudes are consistent with passively evolving L^* galaxies at $z = 0.734$ with $z_{\text{formation}} > 2$ (Bremer 1997).

3C 280 is an FR II radio galaxy with $z = 0.996$ (Spinrad et al. 1985) at $12^{\text{h}}56^{\text{m}}57^{\text{s}}.1$, $+47^{\circ}20'20''$ (J2000.0). It is one of the most powerful FR II galaxies known at radio frequencies. It has an extended ($11''$) emission-line nebula ([O II] $\lambda 3727$), with a central peak and a loop around the eastern radio lobe (McCarthy et al. 1987).

The Best (2000) *K*-band images of both of these sources suggest a richness class 0 environment. At low redshift, richness class 0 clusters of galaxies tend to have X-ray luminosities between 10^{43} and 10^{44} ergs s^{-1} .

ROSAT observers reported extended emission from both of these sources (Hardcastle & Worrall 1999; Crawford et al. 1999; Worrall et al. 1994). For 3C 254, residuals with extent of $30''$ – $60''$ were estimated from PSPC images (Hardcastle & Worrall 1999). Crawford et al. (1999) identify excess emission out to at least $15''$ (145 kpc h_{50}^{-1}) in 3C 254, with 12%–19% of the total X-ray luminosity in extended emission, corresponding to $L_X \sim (5$ – $9) \times 10^{44} h_{50}^{-2}$ ergs s^{-1} . For 3C 280, Worrall et al. (1994) reported PSPC evidence for a cluster extent of $18''$ – $65''$ (200 – $700 h_{50}$ kpc). Hardcastle & Worrall (1999) reported rest-frame 2–10 keV X-ray luminosities of 1.3×10^{45} and $2.8 \times 10^{44} h_{50}^{-2}$ ergs s^{-1} for 3C 254 and 3C 280, respectively. Correspondingly, Hardcastle & Worrall (2000) estimated rough ICM pressures of $\sim 7 \times 10^{-12}$ ($\sim 8 \times 10^{-13}$) ergs cm^{-3} for 3C 254 (3C 280), assuming 7.7 (5.0) keV gas temperatures. In § 5.3, we have reanalyzed these *ROSAT* data for comparison purposes.

3. CHANDRA OBSERVATIONS AND ANALYSIS

Both 3C 280 and 3C 254 were observed with the *Chandra* ACIS-S in VFaint data mode and TIMed read mode. We reprocessed the level 1 events file using updated gain maps from CALDB, Version 2.9 (2001-10-22), and CIAO,

Version 2.2.³ Only events with the *ASCA* grades of 0, 2, 3, 4, and 6 and a status = 0 (“clean”) were used in the subsequent analyses.

3C 254 was observed for 29,668 ks (all corrections applied) on 2001 March 26. The observation was uncontaminated by flare activity, so the good time intervals, as supplied by the original pipeline, were used to extract the final clean data.

3C 280 was observed for 63,528 s (standard good time intervals) on 2001 August 27. A light curve of the background light between 0.3 and 8.0 keV reveals several flares. Even the “quiescent rate” for the observations varied several times. For analysis of flare-free clean data, we only use the 18,944 s of data near the beginning of the observation, particularly for the analysis of extended sources. The signal-to-noise ratio (S/N) estimates of point sources in these clean data are extremely compromised relative to the total, so for analysis of the point sources, which are relatively unaffected by background (usually $\lesssim 10\%$ of the total count rates), we also analyzed the full data set for 3C 280. For the point sources, the count rates and spectra from the full data set are statistically consistent with those extracted from the clean data and, therefore, for the spectral fits (only for the purposes of improving flux estimates of the point sources and for the approximation of their broadband spectra), were done with the full data set of 63,528 s.

For 3C 254, we locate the position of a powerful central X-ray source, confirming the very short projected distance between the nuclear source and the western radio lobe seen by Owen & Puschell (1984). The eastern radio hot spot was not detected (it may be too close to the central source), but the western hot spot has a faint X-ray counterpart (Fig. 1). For 3C 280, we detect both the western and eastern hot spots as well as the nuclear source in the X-ray (Fig. 2). There may be faint emission associated with the bridge; however, estimates of contamination from the scattering wings of the three point sources show this emission to be very uncertain, whether evaluated in the restricted or the full data set. Neither 3C 254 nor 3C 280 (using only the cleaned data set) showed any evidence for the putative cluster sources expected from the previous analyses of the *ROSAT* data for these same sources. Discussion of the *Chandra* limits are in § 5.2.

For the nuclei and the hot spot sources, we used a 4 pixel (1 pixel = $0''.4920$) radius aperture to extract counts. A 4 pixel radius encircles approximately 90% of the energy at 1.7 keV, a correction that is not included in the fluxes and luminosities reported in Table 1. We also extracted a count estimate for the bridge region of 3C 280, excluding the three point sources, for a region of 35×10 pixels ($17''.22 \times 9''.42$), elongated exactly east-to-west, with three $r = 4$ pixel regions excluded for a total area of 150 arcsec². The background-subtracted count rates (0.3–10.0 keV) inside those apertures and the percentages of the total count rates in background counts are reported in Table 1. For each source, a PI spectrum was extracted and binned to at least 15 counts per PI bin. We used XSPEC, Version 11.1, to fit a power-law spectrum to the 0.3–8.0 keV data, with Galactic absorption, fixed at 1.2×10^{20} cm⁻² for 3C 280 and 1.8×10^{20} cm⁻² for

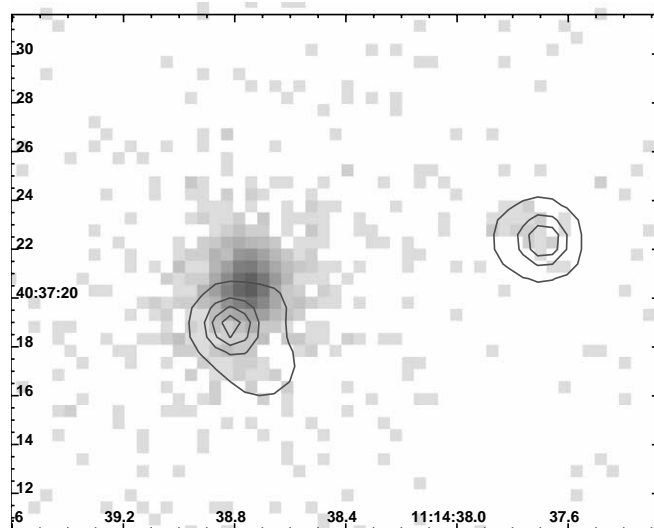


FIG. 1.—Radio contours overlaid on the X-ray gray-scale images for 3C 254 and 3C 280. The 3C 254 radio intensity map was made with the VLA by Liu et al. (1992), observed on 1989 March 21, with a resolution of $1''.25 \times 1''.25$ at 4885.10 MHz. [See the electronic edition of the *Journal* for a color version of this figure.]

3C 254 (Dickey & Lockman 1990; Stark et al. 1992). The fit range is approximate because of the width of the bins in each spectrum. A power law absorbed by the expected Galactic absorption adequately described the spectrum for all of these sources in terms of χ^2 . A thermal spectrum could also be fitted to most of the spectra, except for those of the central sources of 3C 254 and 3C 280. The central source of 3C 280 seems to be unusually flat, with a best-fit spectral index near 0.0; the cleaned data were consistent with that result, with a truncation near 8 keV, but the poor statistics of those flare-affected data limit our ability to confirm the power-law index from the full data set—the data from the clean set are consistent with the fit from the full set but cannot be used for an independent measurement of the slope of the power law.

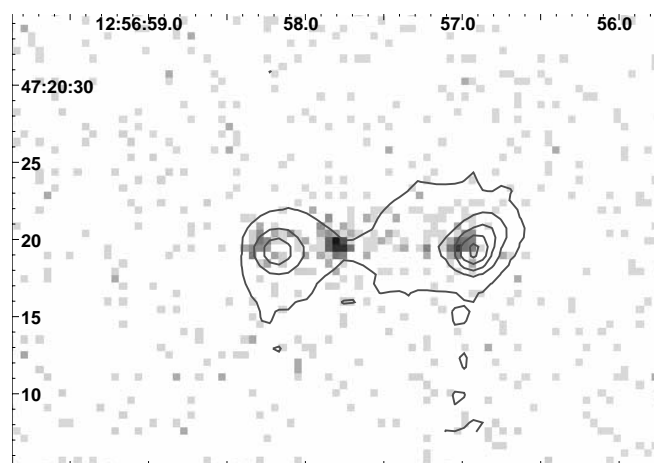


FIG. 2.—Radio contours overlaid on the inverse gray-scale image of the *Chandra* X-ray image for 3C 280. The 3C 280 radio map was obtained from Liu et al. (1992). The radio observation was made with the VLA on 1989 January 21 at 1464.90 MHz, with $1''.20 \times 1''.20$ resolution. [See the electronic edition of the *Journal* for a color version of this figure.]

³ CIAO is the *Chandra* Interactive Analysis of Observations. Version 2.2.1 was released on 2001 December 13 and is available at <http://cxc.harvard.edu/ciao>.

TABLE 1
COMPACT X-RAY SOURCE RESULTS

Feature	Count Rate ^a ($\times 10^{-3}$ counts s ⁻¹)	Flux (0.5–2.0 keV) ($\times 10^{-15}$ ergs s ⁻¹ cm ⁻²)	Power-Law Index	f_v (1 keV) (nJy)	L_X (2–10 keV) ^b ($\times 10^{43}$ ergs s ⁻¹)
3C 280 nucleus	1.94 ± 0.17 (96%)	3.3	0.0 ± 0.36	0.46	14
3C 280 western hot spot	1.07 ± 0.13 (93%)	2.4	2.3 ± 1.0	0.79	2.0
		2.3	$2.2^{+1.4}_{-1.0}$ keV ^c		
3C 280 eastern hot spot	0.60 ± 0.10 (89%)	1.1	2.2^d	0.34	1.2
		1.1	2.0^d		
3C 280 bridge ^e	0.55 ± 0.12 (67%)	2.1	2.6 ± 1.4	<0.60	<1.5
		2.3	$1.3^{+0.2}_{-0.6}$ keV ^c		
3C 280 background sources ^f	3.17 ± 0.23 (90%)	7.1	2.2		
3C 254 western hot spot	0.84 ± 0.19 (81%)	1.4	$2.0^{+0.9}_{-0.7}$	0.52	0.86
		1.9	$2.6^{+1.7}_{-1.5}$ keV ^c		
3C 254 nucleus	139 ± 2 (100%)	311	1.8 ± 0.1	309	320

^a Count rates (background-subtracted) and background were for 0.3–10.0 keV limits. The rates quoted for 3C 280 are for the full data set. Background rates were computed locally to the source; the percentage of the total counts owing to the source is reported in parentheses.

^b Rest frame X-ray luminosity for $H_0 = 50$ km s⁻¹ Mpc⁻¹, $q_0 = 0.0$.

^c kT value.

^d Insufficient counts were available for a fit, so a power-law index of 2.2 was assumed to estimate L_X .

^e Excess counts populate the bridge region of 3C 280, but since there is also contamination from the scattering halos surrounding the hot spots and the nucleus, the detection is marginally significant. To be conservative, we treat the detection as an upper limit.

^f Five compact background sources separated from the nucleus by less than $120''$. Sources were extracted and binned into a single spectrum to obtain an approximate total count rate and flux.

The central source of 3C 254 was affected by pileup at about the 10% level. Pileup flattened the observed spectrum and produced a broad “emission” feature at around 2 keV. We first fitted the spectrum without a pileup model to obtain a best first-guess estimate of the slope and normalization of $\alpha = 1.3$ and 1.4×10^{-4} , respectively (the normalization units are photons keV⁻¹ cm⁻² s⁻¹ at 1 keV). The inclusion of the pileup correction in XSPEC induces unstable fitting in local minima unless a reasonable first guess is provided. With the XSPEC pileup model, the frame time set to 3.2 s, the maximum number of piled-up photons set to 5, $\alpha_{\text{pileup}} = 0.5$, and $g0 = 1$ (pileup model parameters fixed at recommended values), the best-fit slope (photon index) was $\alpha = 1.8 \pm 0.1$, $N_H = 3.2 \times 10^{20}$ (1.8–6.2) cm⁻², and normalization was 4.9×10^{-4} . The reduced $\chi^2 = 1.2$ for this fit with 161 degrees of freedom obtains a null hypothesis probability of 0.05. This power-law index somewhat exceeds the best-fit power-law index from *ASCA* data of $\alpha = 1.67 \pm 0.04$ obtained by Sambruna, Eracleous, & Mushotzky (1999), but the estimates overlap at the 90% uncertainty level, which, given the uncertainties in the pileup model, is not bad.

For 3C 280 and 3C 254, thermal spectra from a gas with $kT \sim 2$ keV and metallicity of $\sim 30\%$ solar also adequately described the spectrum for the hot spot sources. The excess counts from the bridge of 3C 280 could be described by a cooler thermal spectrum of $1.3^{+0.2}_{-0.6}$ keV (with no constraints on the metallicity). As noted earlier, this source is likely to be heavily contaminated by scattered light from the hot spots and the nucleus, so its flux is highly uncertain. Also, since it is not a point source, its spectral properties can be affected by how the background is handled. The inferred broadband fluxes and luminosities for all the sources depend somewhat on the assumption of a spectral model, whether power law or thermal, so we report those properties under both model assumptions in Table 1 for those sources with insufficient signal to distinguish between the two spectral models.

4. HUBBLE SPACE TELESCOPE OBSERVATIONS AND MEASUREMENTS

In order to detect or place limits on optical synchrotron emission from the radio hot spots, we retrieved *HST* Wide Field Planetary Camera 2 (WFPC2) data from the *HST* Archive for both sources. We used co-added and cosmic-ray rejected data from both the CADC and the ECF *HST* data archives and compared photometry from those co-added data with that from individual observations retrieved from the *HST* data archive. The co-addition process (with averaging) obtained consistent photometry for these observations. The gain for all observations was 7.0 electrons DN⁻¹ (data number). The observations were reduced with standard calibration software using the most recent calibration files, switches, and software. No astrometric shifts were applied to the *HST* data. The IRAF/STSDAS task “invmetric” was used to derive the x - and y -positions corresponding to the right ascension and declination of the X-ray hot spots.

The target 3C 254 was observed in a Planetary Camera (PC) snapshot program in both the F702W (R -band) and the F555W (V -band) filters. The nucleus was saturated, but the host galaxy was detected (Lehnert et al. 1999). We do not detect a counterpart for the radio hot spots in either image. We report a 3σ upper limit for a $1'' \times 1''$ aperture centered on the X-ray position of the western hot spot for the F702W observation in Table 2.

The target 3C 280 was observed in the WFPC2, Wide Field 3, for a total of 8800 s in the F622W filter. Ridgway & Stockton (1997) report the unusual appearance of the central galaxy. We find two point-source features within $0''.67$ and $0''.45$ of the western and eastern hot spot *Chandra* positions and aligned with the peaks of the radio hot spots in the Liu, Pooley, & Riley (1992) maps (Fig. 3). We extracted background-subtracted aperture ($r = 0''.3$) photometry at the positions of the *HST* point sources. Using the basic flux calibration information from the header (the PHOTFLAM

TABLE 2
SUMMARY OF *HST* OBSERVATIONS AND RESULTS

Target	3C 254 Western Hot Spot	3C 280 Western Hot Spot	3C 280 Eastern Hot Spot
<i>Chandra</i> position R.A. (J2000.0)	11 14 37.7	12 56 57.04	12 56 58.3
<i>Chandra</i> position decl. (J2000.0).....	40 37 22.4	47 20 19.4	47 20 19
Radio position R.A. (J2000.0).....	11 14 37.7	12 56 56.93	12 56 58.18
Radio position decl. (J2000.0).....	40 37 22.3	47 20 19.3	47 20 19.2
<i>HST</i> position R.A. (J2000.0).....	Not detected	12 56 56.96	12 56 58.26
<i>HST</i> position decl. (J2000.0).....	Not detected	47 20 19.9	47 20 19.3
<i>HST</i> Program ID	5476	5401	5401
Observation date.....	1995 Jan 1	1994 Aug 22	1994 Aug 22
<i>HST</i> filter (WFPC2 chip)	F702W (PC)	F622W (WF3)	F622W (WF3)
Total <i>HST</i> exposure (s).....	2×140	8×1100	8×1100
<i>HST</i> central wavelength (Å).....	6918	6190	6190
PHOTFLAM keyword [ergs s ⁻¹ cm ⁻² Å ⁻¹ (DN s ⁻¹) ⁻¹].....	1.88E-18	2.81E-18	2.81E-18
Aperture (arcsec).....	1.0×1.0	$r = 0.3$	$r = 0.3$
Flux (nJy)	<8700 ^a	985 ± 25	230 ± 15

NOTE.—Units of right ascension are hours, minutes, and seconds, and units of declination are degrees, arcminutes, and arcseconds. Uncertainties are 1σ . Fluxes are corrected for geometric effects (1%), CTI (2%–3%), and aperture corrected for a 0".3 aperture to an infinite aperture (16%).

^a 3σ upper limit, uncorrected for CTI or geometry ($\leq 10\%$ effects).

keyword), we converted the observed count rates (in units of DN s⁻¹) to flux densities in nJy at 6200 Å. We corrected these fluxes for geometric distortion ($\sim 1\%$; Casertano & Wiggs 2001⁴), charge transfer inefficiency ($\sim 2\%$ – 3%), and

aperture ($\sim 16\%$; Whitmore & Heyer 2002;⁵ Mobasher 2002).

The optical flux densities of these point sources are consistent with optical synchrotron emission from the radio

⁴ Casertano & Wiggs (2001, STScI Instrument Science Report, WFPC2 2001-10) is available at http://www.stsci.edu/instruments/wfpc2/Wfpc2_isr/wfpc2_isr0110.html.

⁵ Whitmore & Heyer (2002, WFPC2 Instrument Science Report, 2002-03) is available at http://www.stsci.edu/instruments/wfpc2/Wfpc2_isr/wfpc2_isr0203.html.

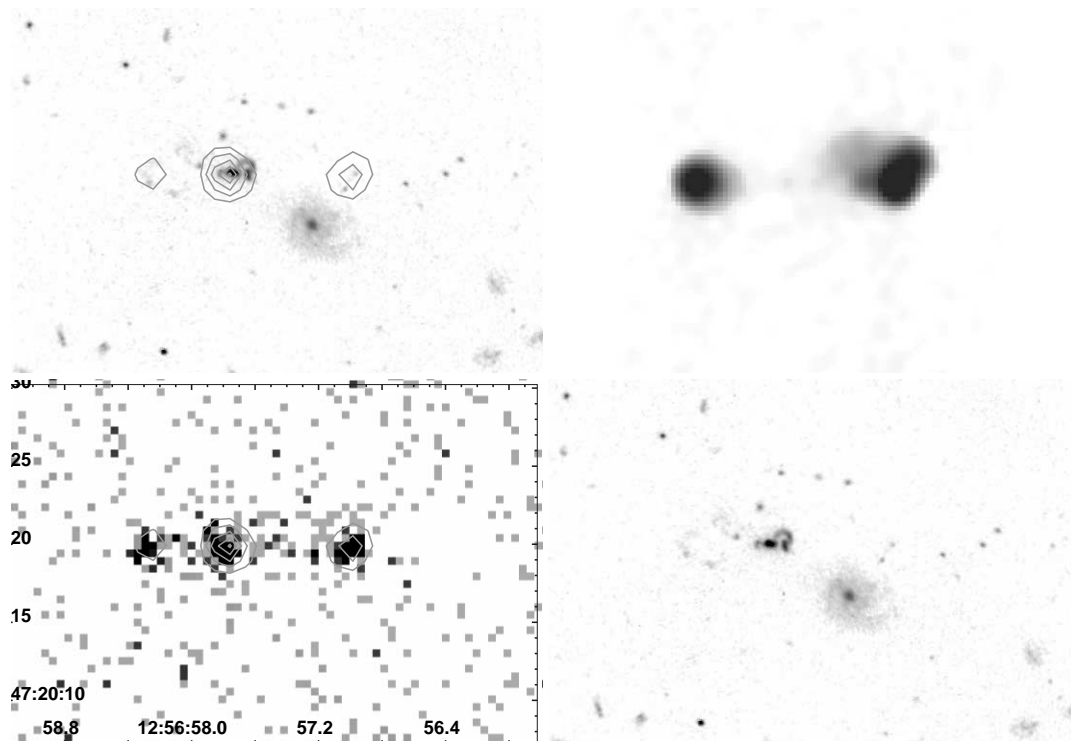


FIG. 3.—Four observations of 3C 280 in the optical, radio, and X-ray, displayed with the same scale and orientation. The upper left-hand figure is the *HST* WFPC2 image of 3C 280 with X-ray contours overlaid (sufficient to show the location of the core and the two hot spots); the lower right-hand figure is the same gray-scale image with the X-ray contours omitted. The upper right-hand figure is the VLA radio image from Liu et al. (1992); the lower left-hand figure is the X-ray image with the same X-ray contours as the upper left *HST* image. We frame the lower left-hand figure with the coordinate system of all four images. The nucleus of 3C 280 seems to be coincident with feature “a” from Ridgway & Stockton (1997). [See the electronic edition of the *Journal* for a color version of this figure.]

and X-ray hot spots but are very unlikely to be produced by SSC or IC of the cosmic microwave background (CMB). We will discuss this topic further in § 6.

5. EVALUATION OF EXTENDED X-RAY EMISSION

The distribution of X-ray emission in clusters of galaxies is typically extended, with a core radius of $\sim 100\text{--}200 h^{-1}$ kpc, and is centered near the center of mass of the cluster. The X-rays arise from an optically thin plasma confined by the gravitational potential of the cluster; the temperature of the gas is indicative of the depth of the potential. Cluster luminosities are typically $10^{43}\text{--}10^{45}$ ergs s^{-1} . Cluster temperatures run between 2 and 15 keV, and typical central densities range between 10^{-2} and 10^{-4} cm^{-3} . In the following section, we describe the analysis performed in order to place constraints on the presence of an extended component arising from a cluster of galaxies centered on these radio sources.

5.1. Radial Profile of 3C 254

A visual inspection of the 3C 254 X-ray image reveals a faint halo around the central bright source. Here we show that this halo is consistent with being part of the point-spread function (PSF) of a very bright *Chandra* X-ray source.

We constructed a model PSF from the library of *Chandra* PSFs by extracting PSFs for the on-axis position on the ACIS-S from the library for six different energies: 0.80, 1.25, 2.25, 3.50, 4.50, and 5.50 keV. Based on the energies of photons in the central X-ray source of 3C 254, each PSF was weighted to construct a summed, normalized PSF. In Figure 4, the resulting radial profile of the predicted PSF is plotted against the radial profile of the central source, extracted identically. Inspection of this plot reveals that there is no detectable excess within $10''$ of the central source over the halo of a bright point source. We therefore exclude a region of radius $12''$ around the X-ray-bright nuclear

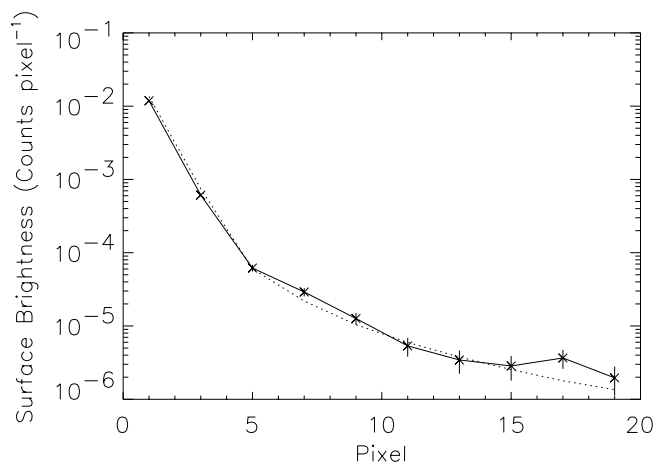


FIG. 4.—Radial profile of the relative surface brightness, in counts per detector pixel ($\sim 0''.5 \times 0''.5$) per second, normalized, of the central source of 3C 254 (solid line) and a radial profile of the weighted distribution from summed PSFs from the *Chandra* PSF library (dotted line). No significant excess is seen in the central $10''$ (20 pixels).

source of 3C 254 when placing a limit on cluster emission around this radio quasar (§ 5.2).

5.2. Upper Limits to Cluster Emission and to Density Profiles of Hot Gas

We used SHERPA, CIAO's fitting engine, to fit the extended, flare-filtered data within $\sim 3 h_{50}^{-1}$ Mpc for both targets between 0.3 and 2.0 keV with the point sources detected by the CIAO routine “wavdetect” masked out (see Fig. 5). For 3C 254, we also masked out a larger region ($\sim 12''$) around 3C 254 and a region around the readout streak. The 0.3–2.0 keV energy range maximized our sensitivity to soft thermal emission while minimizing the contribution from the particle backgrounds. Images were binned by a factor of 4 ($\sim 2''$ pixels) to increase the S/N of each pixel and to reduce the computation time for fitting. Using the Cash

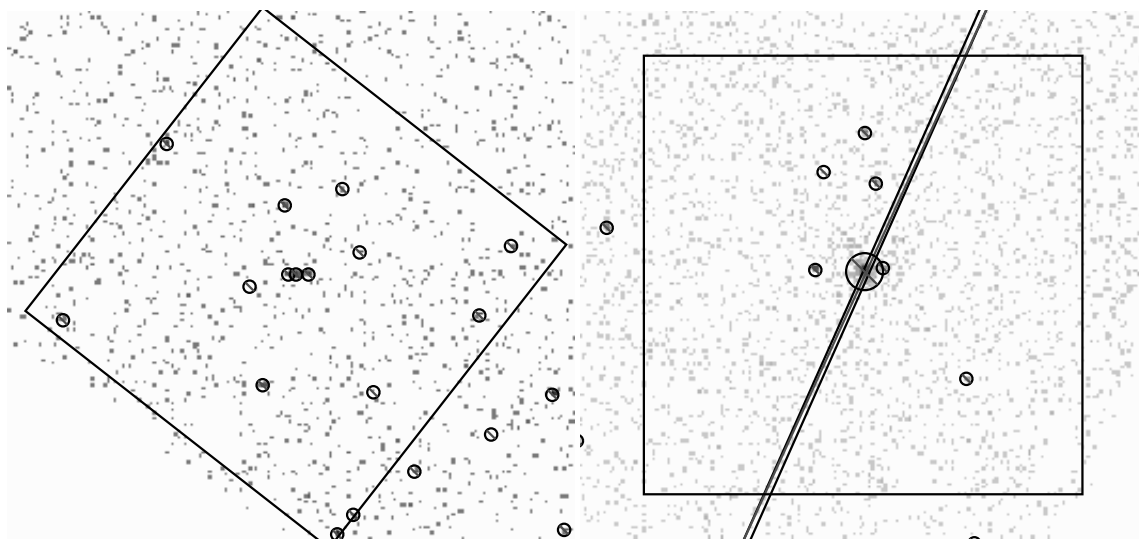


FIG. 5.—Image of a portion of the *Chandra* ACIS-S3 chip (binned by 4) for 3C 280 (left) and 3C 254 (right). In each image, the large box shows the region used in surface brightness fitting. The circles with slashes show the excluded regions (usually sources found by “wavdetect”). For 3C 254, the diagonal box shows the region excluded because of the readout streak.

TABLE 3

3 σ UPPER LIMITS ON CENTRAL SURFACE BRIGHTNESS FOR 29 KS IMAGE OF 3C 254 IN THE 0.3–2.0 KEV BAND

r_c (h_{50}^{-1} kpc)	β -VALUE (counts s $^{-1}$ arcmin $^{-2}$)				
	0.55	0.60	0.65	0.70	0.75
150.....	2.13E-02	2.54E-02	3.00E-02	3.51E-02	4.06E-02
200.....	1.40E-02	1.61E-02	1.84E-02	2.09E-02	2.35E-02
250.....	1.04E-02	1.17E-02	1.31E-02	1.46E-02	1.62E-02
300.....	8.39E-03	9.25E-03	1.02E-02	1.12E-02	1.22E-02
350.....	7.09E-03	7.69E-03	8.36E-03	9.08E-03	9.83E-03

statistic to assess the relative goodness of fit, we fitted a circular β -model profile to the observed surface brightness S_X with the function $S_X = S_0[1 + (\theta/\theta_C)^2]^{-3\beta+1/2}$ and a constant background. S_0 is the normalization, θ is the angular separation from a center (assumed to be aligned with the central X-ray point source), θ_C is the core radius, and β is the slope of the traditional cluster β -function (typically $\sim 2/3$). The β -model slope β and the core radii (θ_C) we tested span a range of standard cluster and group values. We then derived the 3 σ upper limit to the best-fit central surface brightness using the SHERPA “projection” command (Tables 3 and 4) for each pair of β - θ_C values. To infer a limit on the cluster luminosity and central density, we derive the limit on the total count rate within an aperture of radius $2 h_{50}^{-1}$ Mpc, then convert that to a bolometric luminosity limit by assuming a temperature within 1.0 keV of the Markevitch (1998) L_X - T_X relation. Approximate limits on L_{bol} are $2.5 \times 10^{43} h_{50}^{-2}$ ergs s $^{-1}$ for $N_H = 1.2 \times 10^{20}$ cm $^{-2}$ and 1.87×10^{20} cm $^{-2}$ (Dickey & Lockman 1990; Stark et al. 1992) for 3C 280 and 3C 254 (Table 5).

We convert the central surface brightness to a central electron number density, following the derivations for the emission measure from equation (5.68) in Sarazin (1988), where $n_e/n_p = 1.209$ for a fully ionized plasma of solar abundances. For a spherically symmetric plasma, the central electron density in units of cm $^{-3}$ is

$$n_e = \sqrt{\frac{(n_e/n_p)4\pi^{1/2}k(1+z)^2S_0\Gamma(3\beta)}{3.0856 \times 10^7 r_c \Gamma(3\beta - 1/2)}}, \quad (1)$$

where r_c (core radius) is in kiloparsecs. We define k here to be the XSPEC normalization of a Raymond-Smith (or MEKAL) thermal spectrum divided by the *Chandra* count rate in the source spectrum. In this definition, $\int n_e n_H dV = k\text{CR} \times 10^{-14}/4\pi[D_A(1+z)]^2$, where CR is the total count rate (in units of counts s $^{-1}$) from the volume dV , D_A is the angular size distance to the source (in centimeters), n_e is the electron density (in units of cm $^{-3}$), n_H is the hydrogen density (in units of cm $^{-3}$), and V is volume (in units of cm 3 ; XSPEC, Version 11.1; Raymond-Smith model). The

TABLE 4

3 σ UPPER LIMITS ON CENTRAL SURFACE BRIGHTNESS FOR 19 KS IMAGE OF 3C 280 IN THE 0.3–2.0 KEV BAND

r_c (h_{50}^{-1} kpc)	β -VALUE (counts s $^{-1}$ arcmin $^{-2}$)				
	0.55	0.60	0.65	0.70	0.75
150.....	8.09E-03	9.49E-03	1.10E-02	1.26E-02	1.42E-02
200.....	5.44E-03	6.21E-03	7.05E-03	7.92E-03	8.84E-03
250.....	4.09E-03	4.57E-03	5.10E-03	5.66E-03	6.24E-03
300.....	3.29E-03	3.62E-03	3.98E-03	4.36E-03	4.76E-03
350.....	2.78E-03	3.01E-03	3.26E-03	3.53E-03	3.82E-03

normalization per unit count rate k depends on bandpass, temperature, metallicity, and absorption column. The quantity S_0 is in units of counts s $^{-1}$ sr $^{-1}$. We compute the total *projected* emission measure along our line of sight to the cluster inside an aperture of $2 h_{50}^{-1}$ Mpc, derived assuming Galactic absorption and Raymond-Smith models in XSPEC with the response matrices appropriate to the center of the ACIS-S chip.

The normalization k for the 0.3–2.0 keV count rates from a 1.0 keV plasma with a 30% solar metal abundance (typical of a cluster of galaxies at $z < 0.8$; Donahue et al. 1999) is 1.43×10^{-2} in units of XSPEC normalization per *Chandra* count rate (counts s $^{-1}$) for 3C 280 and 1.08×10^{-2} for 3C 254. The conversion factor from central surface brightness (in units of counts s $^{-1}$ arcmin $^{-2}$) to bolometric luminosity for the same plasma, $\beta = 0.7$ and $r_c = 250 h_{50}^{-1}$ kpc, is 5.02×10^{46} and $2.28 \times 10^{46} h_{50}^{-2}$ ergs s $^{-1}$ (counts s $^{-1}$ arcmin $^{-2}$) $^{-1}$ for 3C 280 and 3C 254, respectively.

The 3 σ upper limits on the central electron density, given β , r_{core} , and $k(T_X, Z_\odot, N_H)$, are derived using equation (1). In Table 5, we report 3 σ limits on the bolometric X-ray luminosity (L_{bol}) inside $2000 h_{50}^{-1}$ kpc, the central electron density (n_e), and the electron density [$n_e(r)$] at the hot spot radii r , given $r_{\text{core}} = 250 h_{50}^{-1}$ kpc and $\beta \sim 0.7$ (typical of clusters of galaxies). We estimate $n_e(r)$ to compare to the electron density $n_w(r)$ computed by WDW97. We also explored the same limits for a range of $\beta = 0.55$ – 0.75 and core radius $r_{\text{core}} = 150$ – $350 h_{50}^{-1}$ kpc. All combinations of those two parameters predicted central electron densities of $\sim (2$ – $14) \times 10^{-4} h_{50}^{1/2}$ and not much smaller densities at $78 h_{50}^{-1}$ kpc (since $r < r_{\text{core}}$), where the WDW97 predictions were made based on the radio sources.

WDW97 predict two sets of ambient gas densities: one set assuming minimum-energy conditions to obtain the magnetic field strength in the radio bridge region and one set assuming that the magnetic field strength in the radio bridge is about $\frac{1}{4}$ of the minimum-energy value, as suggested by comparisons of the radio and X-ray-determined ambient gas densities around 3C 295 (Perley & Taylor 1991) and Cyg A (Carilli et al. 1991). This magnetic field strength is

TABLE 5

3 σ UPPER LIMITS ON X-RAY LUMINOSITIES AND ELECTRON DENSITIES

Name	Central Rate (counts s $^{-1}$ arcmin $^{-2}$)	L_{bol} (ergs s $^{-1}$)	kT_{est} (keV)	n_e ($h_{50}^{1/2}$ cm $^{-3}$)	r (h_{50}^{-1} kpc)	$n_e(r)$ ($h_{50}^{1/2}$ cm $^{-3}$)	n_w^{pred} ($h_{50}^{20/7}$ cm $^{-3}$)
3C 254...	1.46E-2	3.4E43	1.5	9.15E-4	78	8.3E-4	$(2.1 \pm 1.0) \times 10^{-4}$
3C 280...	5.66E-3	3.6E43	1.6	9.60E-4	76	8.7E-4	$(8.3 \pm 2.8) \times 10^{-4}$

parameterized by $b = B/B_{\min}$. WDW97 assumed a value of Hubble's constant of $100 \text{ km s}^{-1} \text{ Mpc}^{-1}$. For the case $b = 1$, synchrotron cooling dominates over IC cooling with the CMB radiation, and the radio-determined ambient gas density scales as $H_0^{12/7}$ (see the footnote to Table 4 of Guerra, Daly, & Wan 2000), so for the value of H_0 adopted here, all of the ambient gas densities listed in Table 1, column (9), of WDW97 decrease by a factor of 0.3. Thus, the ambient gas density a distance of about $78 h_{50}^{-1} \text{ kpc}$ from the nucleus of 3C 254 is predicted to be about $1.5 \times 10^{-3} \text{ cm}^{-3}$. The ambient gas density for the western lobe of 3C 280 at a distance of about $76 h_{50}^{-1} \text{ kpc}$ from the center of the radio source is predicted to be about $7 \times 10^{-3} \text{ cm}^{-3}$, while that for the eastern lobe at a distance of about 62 kpc from the radio source center is about $2.4 \times 10^{-3} \text{ cm}^{-3}$ for $b = 1$. These ambient gas densities are significantly larger than the 3σ upper bounds provided by the *Chandra* data for both sources. For n_W for 3C 254, $b = 1$ estimates of n_W could be consistent with the X-ray upper limits if $r_c < 150 h_{50}^{-1}$ and $\beta > 0.7$, values that are more consistent with belonging to groups than clusters.

For the case $b = \frac{1}{4}$, the radio-predicted ambient gas densities are consistent with the X-ray bounds for most values of β and r_c . We report the converted densities (n_W) in Table 5. For $b = \frac{1}{4}$, the bridge magnetic energy density for the bridges of 3C 254 and 3C 280 is comparable to the energy density of the CMB, so IC cooling is as important in these regions as synchrotron cooling, and the radio-predicted ambient gas density scales as $H_0^{20/7}$ (see the footnote to Table 4 of Guerra et al. 2000). For the case $b = \frac{1}{4}$, the predicted ambient gas densities are obtained from column (11) of Table 1 of WDW97 and multiplied by the factor of $0.14 = 0.5^{20/7}$ to account for the value of H_0 adopted here. The radio-predicted ambient gas densities are then about $2.1 \times 10^{-4} \text{ cm}^{-3}$ for the western lobe of 3C 254 and about $8.3 \times 10^{-4} \text{ cm}^{-3} h_{50}^{20/7}$ for both the western and eastern lobes of 3C 280. These gas densities predicted by WDW97 (n_W) for $b = \frac{1}{4}$ are listed in Table 5. For both sources, the radio-predicted ambient gas density is consistent with the upper bound provided by *Chandra* for $b = \frac{1}{4}$.

The radio-predicted ambient gas density has a very strong dependence on H_0 and thus also the coordinate (or proper) distance ($a_0 r$) to the source. When IC cooling with CMB photons is important, the ratio of the radio-determined ambient gas density n_W to the X-ray-determined value n_e is $n_W/n_e \propto (a_0 r)^{-2.36} \propto h_{50}^{2.36}$. If we shift from the cosmology assumed here, $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $q_0 = 0$, to a cosmology with $H_0 = 65 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.3$, and $\Omega_\Lambda = 0.7$, the ratio between the predicted value by WDW97 and the 3σ upper limit from the *Chandra* observations is increased by a factor of about $0.8^{-2.36} \simeq 1.7$. This effect does not change our conclusion presented above for 3C 254, but it does affect the comparison between the radio-predicted ambient density and the upper limit on the density from the X-ray observations of 3C 280, where the radio-predicted ambient density may be in conflict with the X-ray observations at about the 1σ confidence level. The radio-predicted ambient gas density is increased by the ratio of the coordinate distances (0.8) to the power 2.86, leading to a predicted ambient gas density for 3C 280 (see Table 5) of about $(16 \pm 5.5) \times 10^{-4} \text{ cm}^{-3}$, whereas the upper bound placed by the *Chandra* observations increases by a factor of the square root of the ratio of the coordinate distances, $0.8^{-0.5}$, to about $10 \times 10^{-4} \text{ cm}^{-3}$.

An additional test could be made using our limits on the temperature of any putative X-ray-emitting gas distribution if the ambient gas pressure based on the radio properties of 3C 254 and 3C 280 could have been determined, as it has for many other high-redshift 3CR sources (see Table 1 of Wan, Daly, & Guerra 2000). The luminosity limits are extremely stringent for both sources, corresponding to bolometric luminosities of $\sim 3.5 \times 10^{43} h_{50}^{-2}$ and temperatures $T_X \sim 1.5 \text{ keV}$.

These observations rule out the presence of a hot, massive cluster with a bright ICM for these particular radio sources. Specifically, we definitively rule out the presence of a cluster with the luminosity and distribution of X-ray plasma like the one containing Cyg A and that of the putative clusters associated with these sources, as inferred from *ROSAT* data (see more in § 5.3).

Our conclusions are consistent with the recent findings from an optical study by Harvanek & Stocke (2002), which found that FR II-selected clusters are less rich (Abell richness class 0–1) than their X-ray-selected cousins, which tend to contain FR I sources exclusively. In their study, based on a review of radio structures and surrounding galactic density for 3CR radio galaxies between $z = 0.15$ and 0.65 , they conclude that radio sources with a large bending angle cannot be used as signposts for rich clusters around powerful radio sources.

5.3. ROSAT PSPC and HRI Fluxes

In this section, we compare the *Chandra* observations with both earlier results from the literature and the archived *ROSAT* observations, in order to try to understand why clusters were reported around these objects. We find that the total fluxes measured by *ROSAT* and *Chandra* within $120''$ are in agreement for 3C 280 and that the nuclear source for 3C 254 has likely varied in the last 5 years. Therefore, at least for 3C 280, the extended emission inferred from the *ROSAT* data is explained by faint point sources.

3C 280 was observed by the *ROSAT* PSPC in 1997 for 47.6 ks. Worrall et al. (1994) found that the PSPC radial profile was best fitted by a point source plus a β -model, although they could not constrain the core radius very well. They stated that 60% of the net counts in the 0.2–1.9 keV band were from an extended component around 3C 280 and a flux density of 1.7 nJy at 1 keV from the extended component. Although not actually given, their total flux density can be computed using quantities in Table 4 of Hardcastle & Worrall (1999). The total flux density at 1 keV is 6.2 (± 3.6) nJy within a source radius of $120''$ (uncertainties were high for the extended component). They also noted a possible asymmetry in the X-ray emission.

Using the same PSPC data, we cleaned three episodes of high background from the data for net 42.7 ks (corrected for dead time). We adaptively smoothed the 3C 280 events with energies between 0.5 and 2.0 keV (the lower limits means we cut out much of the soft X-ray background) to a minimum of 10 counts in the smoothing kernel. Figure 6 shows the resulting PSPC contours overlaid on the 4×4 binned *Chandra* 63 ks image.

The emission has at least three peaks within the $120''$ source radius used by Worrall et al. (1994). The northern peak is definitely seen in the *Chandra* image as a separate point source. Other, fainter sources in the *Chandra* image may also be contributing to the total PSPC flux. Worrall

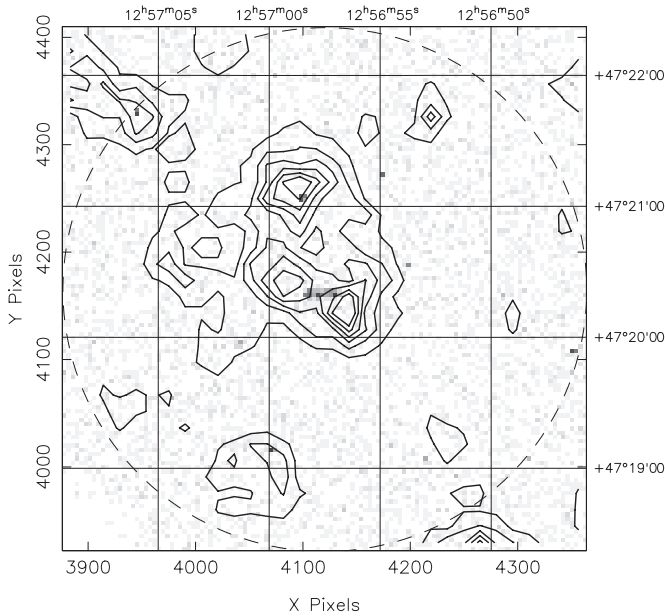


FIG. 6.—*Chandra* ACIS-S3 image (binned by 4) overlaid with *ROSAT* PSPC 0.5–2.0 keV X-ray contours. The contours begin at 1.5 times the background level and increase in steps of 0.5 times the background level thereafter. The dashed circle indicates the 120'' source radius used by Worrall et al. (1994).

et al. (1994) considered all these emission regions as one source.

As a check, we compare the total PSPC flux within a 120'' region around 3C 280 to the summed flux of *Chandra* point sources within the same region. First, we extract the PSPC spectrum for an aperture of 120'' around 3C 280 with a background from an annulus about 3'–4'5, masking out several sources in the background. A custom ARF and standard PSPC RMF were used; the spectrum was binned to eight energy bins. The resulting spectrum was fitted to an $\alpha = 2$ power-law model in XSPEC, Version 11.2, leaving only the model normalization as a free parameter. N_{H} was fixed at the Galactic value (Dickey & Lockman 1990; Stark et al. 1992). We derive a 0.5–2.0 keV flux of $(2.4 \pm 1.5) \times 10^{-14}$ ergs $\text{s}^{-1} \text{cm}^{-2}$, which corresponds to a flux density of 6.6 ± 4.1 nJy at 1 keV. This value is statistically the same as the 6.2 nJy 1 keV flux found by Worrall et al. (1994). A similar analysis of the *ROSAT* HRI data (observed for 53 ks in 1997) gives an HRI count rate of 1.1×10^{-3} counts s^{-1} , which corresponds to a flux of 2.2×10^{-14} ergs $\text{s}^{-1} \text{cm}^{-2}$ or 6.1 nJy at 1 keV.

The sum of the estimated flux densities from *Chandra* sources within 120'' of 3C 280, including five background sources, is $\sim 4.9 \pm 0.6$ nJy (with an approximate 10% correction for the flux outside the 4 pixel apertures). The sum of the point sources and bridge source within 120'' of the 3C 280 nucleus is statistically consistent with, but somewhat lower than, the *ROSAT* PSPC and HRI estimates from the original data. So most, if not all, of the *ROSAT* flux can be accounted for by the point sources in the *Chandra* observation.

3C 254 was observed with the *ROSAT* PSPC for 16 ks in 1997 and with the *ROSAT* HRI for 29 ks in 1996. Hardcastle & Worrall (1999) report a total (extended and pointlike) 1 keV flux density of $\sim 250 \pm 13$ nJy. They found that the radial profiles of both the PSPC and HRI data showed evidence for extended emission, although the HRI data gave a somewhat smaller and less luminous extended component.

We analyzed the PSPC observation of 3C 254 similarly to 3C 280 and, assuming a power-law index of 1.0, found a flux density at 1 keV of 168 ± 10 nJy, somewhat smaller than the Hardcastle & Worrall estimate. Our HRI flux density is also somewhat lower at 201 nJy. However, the *Chandra* estimate of the 1 keV flux density is ~ 300 nJy, larger than either of these values. The small difference in assumed power-law indices (Hardcastle & Worrall 1999 assume a power-law index of 0.8) does not account for the variation. The comparison between the *ROSAT* and *Chandra* fluxes shows what is likely to be long-term X-ray variability of the radio quasar.

Table 6 summarizes our *ROSAT* results.

6. X-RAY EMISSION ASSOCIATED WITH THE RADIO HOT SPOTS

The X-ray emission that is coincident with the radio hot spots could be produced by the SSC, by synchrotron radiation, by IC scattering with external photon fields, by a combination of these processes, or by some other process. SSC and synchrotron radiation are discussed in §§ 6.1 and 6.2. Considerations of IC scattering of an external photon field, such as the CMB radiation, allow a lower limit to be placed on the magnetic field strength of the region, which is discussed in § 6.3. Constraints on the bridge magnetic field strength of 3C 280 are also discussed in this section.

6.1. The Synchrotron Self-Compton Process

Sources with *Chandra* detections of hot spot emission that is most likely due to SSC include 3C 295 (Harris et al. 2000),

TABLE 6
SUMMARY OF *ROSAT* OBSERVATIONS

Source	Instrument	Net Counts ^{a,b}	Radius (arcsec)	Background ^b (counts $\text{s}^{-1} \text{arcmin}^{-2}$)	Exposure Time (ks)	Flux (0.5–2.0 keV) (ergs $\text{s}^{-1} \text{cm}^{-2}$)	F_{ν} (1 keV) (nJy)
3C 280.....	PSPC	58 ± 21	120	$5.00\text{E}-4$	42.7	$2.38\text{E}-14$	6.6
	HRI	41 ± 40	120	$4.44\text{E}-3$	53.3	$2.22\text{E}-14$	6.1
3C 254.....	PSPC	713 ± 29	150	$2.30\text{E}-4$	15.6	$6.12\text{E}-13$	168
	HRI	971 ± 35	150	$4.93\text{E}-3$	29.2	$7.29\text{E}-13$	201

^a Net counts within a circular aperture with radius given in the radius column.

^b For the PSPC, the listed values are in the 0.5–2.0 keV band to eliminate contribution from the soft X-ray background, but for the HRI, they are in the full 0.1–2.4 keV band since the HRI has no spectral resolution.

Cyg A (Wilson, Young, & Shopbell 2000), 3C 123 (Hardcastle, Birkinshaw, & Worrall 2001), and 3C 207 (Brunetti et al. 2002).

SSC emission is most important in bright small regions, such as radio hot spots, while IC upscattering of external photons, such as those that comprise the CMB, is important in the more diffuse lobe and bridge regions of the source. The dominant mechanism in any area is determined by the ratio of the local photon energy densities. IC scattering of local synchrotron photons, known as SSC emission, will dominate over scattering of ambient photons, such as those of the CMB, when the local energy density of synchrotron photons is larger than the energy density of other photon fields. For a uniform spherical synchrotron emitting region with bolometric luminosity L , radius r , and volume V , the local photon energy density is approximately $u_s \approx 0.75Lr/cV$ (Band & Grindlay 1985). The bolometric luminosity $L = \int L_\nu d\nu$ can be written in terms of a bolometric correction factor κ_ν , where $\kappa_\nu = L/L_\nu\nu = f/(f_{\nu_0}\nu_0)$; f is the total flux, $f = \int f_\nu\nu_0$, and f_{ν_0} is the flux density detected at frequency ν_0 . The energy density of synchrotron photons in a region with diameter $2r$ and corresponding angular diameter θ is $u_s \approx 9\kappa_\nu f_{\nu_0}\nu_0(1+z)^4/c\theta^2$. Writing the flux density f_{ν_0} in units of Jy, θ in units of arcseconds, and the frequency ν_0 in units of GHz, the energy density of synchrotron photons of a radio hot spot is

$$u_s \approx 1.3 \times 10^{-13} \kappa_\nu (1+z)^4 f_{\nu_0} \nu_0 \theta^{-2} \text{ ergs cm}^{-3}. \quad (2)$$

A minimum local photon energy density due to external sources is provided by the CMB radiation, which has an energy density of $u_{\text{CMB}} = 4.19 \times 10^{-13} (1+z)^4 \text{ ergs cm}^{-3}$, assuming a zero redshift temperature for the blackbody radiation of 2.728 K (Fixsen et al. 1996). The ratio of the energy densities at the location of the radio hot spot is

$$\frac{u_s}{u_{\text{CMB}}} \approx 0.3 \kappa_\nu \frac{f_{\nu_0}}{\text{Jy}} \frac{\nu_0}{\text{GHz}} \left(\frac{\theta}{\text{arcsec}} \right)^{-2}, \quad (3)$$

which is independent of choices of global cosmological parameters, redshift, and minimum-energy assumptions. Thus, only in small, bright, radio-emitting regions will SSC dominate over IC with CMB photons and other ambient photon fields.

Published radio observations do not resolve the hot spots of either 3C 254 or 3C 280. The highest resolution image of 3C 254 available is a 4.885 GHz image obtained at $0''.5$ resolution (Owen & Puschell 1984). The western hot spot has a flux density of about 0.13 Jy and an angular size less than about $0''.5/(\ln 2)^{1/2}$. Assuming that the angular diameter of the hot spot is less than $0''.5/(\ln 2)^{1/2}$, equation (2) indicates $u_s/u_{\text{CMB}} \gtrsim 1.5$, assuming a value of $\kappa_\nu \sim 3$, typical of sources with confirmed SSC phenomena and observed at similar frequencies, such as 3C 295 (Harris et al. 2000) and Cyg A (Wilson et al. 2000). This energy density ratio indicates that SSC is likely to be more important than IC with external photon fields within the hot spot region. We stress that equation (2) is only a rough approximation obtained assuming that the hot spot is a uniformly emitting sphere whose angular size and radio spectrum can be accurately determined. On a slightly larger scale, about $1''.25$ in diameter, the 1.4 GHz flux density of the western hot spot of 3C 254 is about 0.7 Jy (Liu et al. 1992). Substituting a value of θ of $1''.25/(\ln 2)^{1/2}$ into equation (2) implies that $u_s/u_{\text{CMB}} \sim 0.4$,

suggesting that on scales smaller than $1''$ or so SSC will be more important than IC scattering with CMB photons, whereas on scales larger than an arcsecond, IC scattering with CMB photons will be more important in producing X-rays than SSC.

For 3C 280, the highest resolution radio datum available is the 1.4 GHz map of Liu et al. (1992); the angular resolution is about $1''$. Assuming that the angular diameter of the source is less than about $1''/(\ln 2)^{1/2}$ and adopting $\kappa_\nu \approx 3$, equation (2) indicates that $u_s/u_{\text{CMB}} \gtrsim 1.5$ and 0.7 for the western and eastern hot spots of 3C 280, which have 1.4 GHz radio flux densities of about 2 and 0.9 Jy, respectively. Thus, SSC is likely to dominate over IC scattering within the central $1''$ diameter of the hot spot, while IC scattering with ambient photon fields is likely to dominate outside this region. Both processes may contribute to the X-ray emission, and SSC could be an important process in the hot spot regions of 3C 254 and 3C 280.

A rough estimate of the SSC X-ray flux density using the unresolved radio hot spot data is possible if we make a few assumptions. As discussed by Harris et al. (2000), the ratio of X-ray to radio luminosities depends on the ratio $R = u_s/u_B = L_X/L_s$ when the X-rays are produced by SSC, where u_s is defined above and $u_B = B^2/8\pi$ is the magnetic energy density. The synchrotron photon energy density scales as θ^{-2} (see eq. [1]), and the magnetic energy density scales as $\theta^{-1.7}$ when the magnetic field strength is estimated using minimum-energy assumptions (Burbidge 1956) since $B_{\text{min}} \propto \theta^{-6/7}$. Thus, the ratio of energy densities has a very weak dependence on θ : $R \propto \theta^{-0.3}$ approximately. Since this ratio is so weakly dependent on θ , the ratio determined using the unresolved flux densities predicts L_X/L_s if the observed radio and X-ray emission are produced in the same volume. The X-ray luminosity can be written $L_X = \kappa_X L_{\nu_X} \nu_X$, and the synchrotron luminosity can be written $L_s = \kappa_s L_{\nu_s} \nu_s$. Since the X-ray and radio sources are coincident, $L_{\nu_X} \nu_X / L_{\nu_s} \nu_s = f_{\nu_X} \nu_X / f_{\nu_s} \nu_s$, where $f_{\nu_{X,s}}$ are the observed X-ray and synchrotron flux densities, respectively. Thus, the ratio R can be computed using $R = u_s/u_B$, and the SSC flux density can be predicted using $f_{\nu_X} \approx R f_{\nu_s} \nu_s / \nu_X$, valid for $\kappa_X \approx \kappa_s$. Both the radio and X-ray flux densities are falling with frequency approximately as ν^{-1} , and so both are detected on the declining side of the spectrum; in this case it is expected that $\kappa_X \approx \kappa_s$.

For the western hot spot region of 3C 254, the minimum-energy magnetic field strength is about $180 \mu\text{G}$, and the synchrotron photon energy density is about $6 \times 10^{-12} \text{ ergs cm}^{-3}$ (estimated using the 4.885 GHz data at $0''.5$ resolution), so the ratio of energy densities is about $R \approx 0.005$. This ratio predicts a 1 keV flux density due to SSC of about 0.015 nJy, about a factor of 0.025 of the observed flux density. If the magnetic field strength is reduced by a factor of 6, the observed X-ray flux density matches that observed. It is possible that SSC could account for the observed X-ray emission if the magnetic field strength is approximately $\frac{1}{6}$ of the minimum-energy value. To reliably compute the predicted SSC 1 keV flux density using radio data and to solve for the magnetic field strength in this region, knowledge of the angular diameter of the hot spot is needed in addition to the radio spectrum, including the frequencies of spectral breaks (that is, a full description of the radio spectrum).

For the western hot spot region of 3C 280, the minimum-energy magnetic field is about $180 \mu\text{G}$ and the synchrotron photon energy density is about $10^{-11} \text{ ergs cm}^{-3}$ (obtained

using the 1.4 GHz data at about $1''$ resolution), so the energy density ratio is $R \approx 0.007$ and the predicted 1 keV flux density is about 0.08 nJy. This flux density is about 0.1 that observed, implying a magnetic field strength of $\sim \frac{1}{3}$ the minimum value, if the X-ray emission is produced by SSC. For the eastern hot spot region of 3C 280, the minimum-energy magnetic field is about $170 \mu\text{G}$ and the synchrotron photon energy density is about $4 \times 10^{-12} \text{ ergs cm}^{-3}$, so the ratio $R \approx 0.004$, and the predicted 1 keV flux density is about 0.02 nJy, about 0.06 of that observed. The observed X-ray emission could be produced by SSC if the magnetic field strength is approximately $\frac{1}{4}$ of the minimum-energy value.

In summary, the X-ray emission detected from the hot spot vicinities of 3C 280 and 3C 254 could be produced by SSC if the magnetic field strengths are roughly 0.2–0.3 of the minimum-energy values. Larger magnetic field strengths in and around the radio hot spots relative to the minimum-energy value in that region would produce less X-ray emission via SSC. However, if the angular sizes of the hot spots are substantially smaller than the resolution adopted here or if the value of κ_ν is larger than that adopted here, SSC could produce the observed X-ray emission from the hot spot regions with magnetic field strengths closer to the minimum-energy values.

6.2. X-Ray Synchrotron Radiation

The synchrotron process that produces the radio hot spot emission could extend out to X-ray energies and produce the X-ray hot spots detected by *Chandra*. This process would require quite recent acceleration since the synchrotron aging timescales for very high energy, relativistic electrons that could produce X-rays is quite short (on the order of a human lifetime!). However, since acceleration is almost certainly occurring in the hot spot region, relativistic electrons with very high Lorentz factors could conceivably exist there.

For 3C 280, the 1.4 GHz flux density of western hot spot is about 2 Jy. The 1 keV flux density for the western hot spot is about 0.79 nJy, which would imply a radio-to-X-ray synchrotron spectral index of about 1.14. Similarly, the eastern hot spot 1 keV flux density of 0.34 nJy combined with the 1.4 GHz flux density of about 0.9 Jy also indicates a radio-to-X-ray spectral index of 1.14; these spectral indices are consistent with the 1.4, 5, and 15 GHz radio spectral index determined by Liu et al. (1992). If the synchrotron radiation extends all the way from radio to X-ray energies, the predicted 6200 Å, or 5×10^{14} Hz, flux density is about 935 and 420 nJy for the western and eastern hot spots of 3C 280, respectively.

The hot spot regions of 3C 280 may have optical counterparts emitting 6200 Å radiation detected by *HST* (this paper). The *HST* photometry of point sources nearly coincident with the radio and X-ray hot spots suggests point-source flux densities about 985 ± 25 and 230 ± 15 nJy for the western and eastern hot spots, respectively. These fluxes are certainly in the ballpark of the predicted values and are consistent with synchrotron radiation, particularly if the synchrotron spectrum has some curvature. Considering a broken power law with one spectral index α_{ro} between 1.4 GHz and 6200 Å and a second spectral index α_{ox} between 6200 Å and 1 keV, these spectral indices that describe the data are $\alpha_{\text{ro}} \approx 1.13$ and $\alpha_{\text{ox}} \approx 1.15$ for the western hot spot of 3C 280 and $\alpha_{\text{ro}} \approx 1.19$ and $\alpha_{\text{ox}} \approx 1.05$ for the eastern hot

spot of 3C 280. The indices for the western hot spot seem reasonable, while those for the eastern hot spot would be quite unusual since the high-energy spectral index is expected to be similar to or greater than the low-energy spectral index. Therefore, there may be another source of radiation contributing to the X-ray emission from the eastern hot spot of 3C 280. If the radio-to-optical spectral index of 1.19 is continued from the optical to the X-ray in the eastern hot spot of 3C 280, the 1 keV flux density would be about 0.15 nJy, about half of the detected X-ray flux density (see Fig. 7).

Thus, the data are consistent with synchrotron radiation producing the optical hot spot emission detected by *HST*. The X-ray emission from the western hot spot is consistent with synchrotron radiation, but that from the eastern hot spot would require unusual spectral properties or an additional source of X-ray emission such as SSC or IC scattering of relativistic electrons with CMB photons.

The minimum-energy magnetic fields in the hot spot regions of 3C 280 (on scales of about $1''$) are about $200 \mu\text{G}$. Such a magnetic field strength implies that the 1 keV emission is produced by relativistic electrons with Lorentz factors of about 2×10^7 . This conclusion follows from the fact that relativistic electrons with Lorentz factors of γ in a tangled field, for which the field component perpendicular to the electron velocity (B_\perp) satisfies $B_\perp = (2/3)^{1/2} B$, produce synchrotron radiation with frequency $\nu \approx 30\gamma^2 B_{-5}$ Hz. Here B_{-5} is the magnetic field strength in units of 10^{-5} G. The radiative lifetime of a relativistic electron with a Lorentz factor of 2×10^7 in a magnetic field of about $200 \mu\text{G}$ is only about 30 years (see Daly 1992b, eq. [11b]). This estimate assumes a tangled magnetic field configuration and that pitch angle scattering is important.

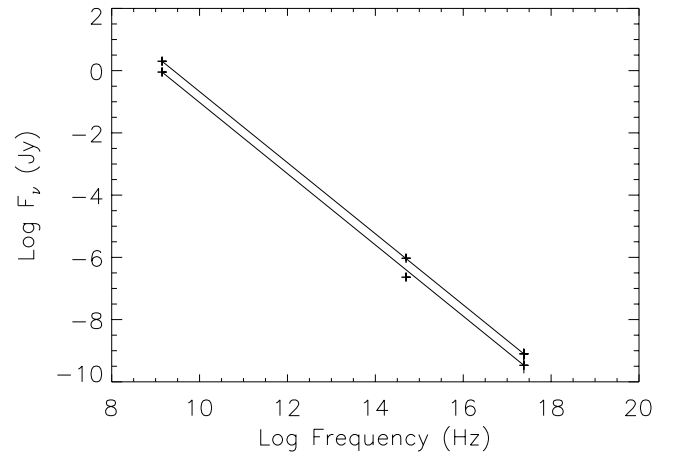


FIG. 7.—Radio, optical, and X-ray flux densities for the western (*upper set of points*) and eastern (*lower set of points*) hot spots in 3C 280. The error bars are 3σ . The radio fluxes have uncertainties of about 10%, while the uncertainties in the optical and X-ray flux densities include a conservative 5% flux calibration uncertainty for both *Chandra* and *HST*. A solid line connects the radio and X-ray points, showing that the optical flux density for the western hot spot falls just where expected if the radio, optical, and X-ray emission are produced by synchrotron radiation. The optical flux density of the eastern hot spot is somewhat below what one might extrapolate for synchrotron emission, based on the radio and X-ray flux densities. Alternatively, one could say that the X-ray flux density in the eastern hot spot is somewhat above what one would expect for synchrotron emission based on the radio and optical flux densities.

For the western hot spot of 3C 254, the 4.885 GHz flux density is about 0.1 Jy. If the 1 keV X-ray emission of ~ 0.5 nJy is synchrotron radiation from the same population of relativistic electrons, the radio-to-X-ray spectral index is about 1.09, and the predicted 6000 Å, or 5×10^{14} Hz, flux density is about 450 nJy. This is well below the current bounds placed on this emission using *HST* data of ~ 8700 nJy (Table 2). Thus, this X-ray emission could also be produced via X-ray synchrotron. The minimum-energy magnetic field on a scale of about $1''$ centered on this hot spot is about $125 \mu\text{G}$, implying that 1 keV synchrotron radiation would be produced by electrons with Lorentz factors of about 2×10^7 , which would have radiative lifetimes of about 70 years.

6.3. Inverse Compton Scattering with the Cosmic Microwave Background

X-ray emission will be produced by IC scattering between relativistic electrons and CMB photons and with other ambient photon fields. Relativistic electrons with Lorentz factors of about 10^3 will IC scatter CMB photons to an observed energy of about 1 keV, independent of any assumptions (including assumptions concerning the magnetic field strength) and independent of redshift (see, e.g., Daly 1992a, 1992b). If the relativistic electron energy distribution does not extend below these values in the hot spot region, adiabatic expansion will cause a shift of the Lorentz factors of the relativistic electrons to lower energies in the regions outside of the hot spots, as discussed in some detail by Daly (1992a). For magnetic field strengths near minimum-energy values, detected radio emission from outside the hot spots will be produced by electrons with Lorentz factors of about 10^3 . Given that the relativistic electron energy distribution extends to these Lorentz factors, which it almost certainly must, a lower limit can be placed on the magnetic field strength of the plasma in the immediate vicinity of the radio hot spot by calculating the 1 keV X-ray flux density of emission produced by IC scattering with the CMB and requiring that this flux density be less than or equal to the observed 1 keV flux density. Using equation (8b) from Daly (1992b) and the observed X-ray-to-radio flux density ratio, the magnetic field strength can be determined or bounded.

For the western hot spot of 3C 254, the 1 keV to 1.4 GHz flux density ratio is about 0.7×10^{-9} , the radio spectral index is about 1.0 (Liu et al. 1992), and the 1.4 GHz radio flux density is about 0.7 Jy. Assuming that the X-ray flux produced by IC with the CMB is less than or equal to that detected, the magnetic field strength in this region must satisfy $B \gtrsim 35 \mu\text{G}$. The minimum-energy magnetic field for the resolved region, obtained using the 1.4 GHz data at about $1''$ resolution and adopting the standard assumptions, is about $125 \mu\text{G}$. This result implies that the magnetic field strength in the hot spot vicinity must be greater than about 0.3 of the minimum-energy value. This field strength relative to the minimum-energy value is slightly larger than that required for the X-ray emission to be completely explained by the SSC process, although given how rough the SSC estimate is, we take the two constraints to be consistent. (It will be improved using higher resolution radio data with broad frequency coverage.) Similar constraints are obtained using the higher resolution radio data of Owen & Puschell (1984) along with the assumption that all

of the X-ray flux originates from a region that is smaller than $0.5''$.

For the western hot spot region of 3C 280, the radio flux density is about 2 Jy at 1.4 GHz at about $1''$ resolution; the radio spectral index of the western hot spot is roughly 0.8, steepening quickly away from the hot spot (Liu et al. 1992). The 1 keV to 1.4 GHz flux density ratio is about 4×10^{-10} . Substituting this value into equation (8b) of Daly (1992b), the magnetic field strength in this region satisfies $B \gtrsim 50 \mu\text{G}$. The minimum-energy magnetic field strength in this region is about $180 \mu\text{G}$, implying that the magnetic field in this region cannot be less than about 0.3 of the minimum-energy value. The eastern hot spot has a 1.4 GHz flux density of about 0.9 Jy, a radio spectral index of about 1, and a 1 keV to 1.4 GHz flux density ratio of about 4×10^{-10} , so the magnetic field strength in this region must be greater than about $65 \mu\text{G}$, or else IC scattering with CMB photons would produce more X-ray emission than is detected. The minimum-energy magnetic field in this region is about $170 \mu\text{G}$, implying that the magnetic field strength in the vicinity of the radio hot spot cannot be less than about 0.4 of the minimum-energy value. Both of these deviations from minimum-energy conditions are quite similar to and consistent with those required for SSC to contribute to the detected X-ray emission.

Thus, the X-ray and radio flux density ratios together with the requirement that the IC scattering between relativistic electrons and CMB photons not produce more X-rays than observed suggest that the magnetic field strengths of the hot spot regions cannot be significantly less than about 0.3–0.4 of the minimum-energy values. These values are similar to those required if the X-ray emission is produced by SSC, although both of these estimates are quite rough. *The similarity of these bounds suggests that both processes could be contributing to the X-ray emission, with SSC being produced in the heart of the hot spots and IC with CMB photons being produced in the more extended hot spot vicinity.*

The fact that a lower magnetic field strength would overproduce X-rays indicates that the relativistic plasma in the radio hot spot cannot depart significantly from equipartition or minimum-energy conditions; the field strength is expected to be on the order of or less than the equipartition field strength, and the X-ray brightness of the hot spots relative to the radio brightness indicates that the field cannot be much less than about 0.3 of the minimum-energy value.

The *Chandra* data also allow an upper bound to be placed on X-ray emission from the bridge region of 3C 280 (see Table 1). This bound can be compared with predicted 1 keV X-ray fluxes produced by IC scattering of microwave background photons with the relativistic electron population in the bridge region. Using AIPS,⁶ we estimated the 1.4 GHz radio flux density within the 150 arcsec^2 region of the radio bridge that matches that used to obtain the X-ray bound to be about 0.4 Jy; this region does *not* include the radio hot spots. Most of this flux density originates from the extremities of the bridge region closest to the radio hot spots (although the hot spot regions are specifically excluded). The minimum-energy magnetic fields in these regions are about $90 \mu\text{G}$ for the western bridge and about $70 \mu\text{G}$ for the

⁶ Astronomical Image Processing System is a software package for reduction and analysis of radio data and is available at <http://www.aoc.nrao.edu/aips>.

eastern bridge (see Table 1 of WDW97, converted using $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$). The radio spectral index in this region is about 1 (Liu et al. 1992). Using equation (8c) of Daly (1992b), the 1 keV upper bound on the X-ray flux can be combined with the radio flux density and spectral index to determine the smallest value of the magnetic field allowed, about $30 \mu\text{G}$. Given the minimum-energy magnetic field strengths listed above, this inferred lower limit means that the magnetic field strength for the bridge regions near the extremities of the radio bridge must be greater than about 0.3–0.4 of the minimum-energy value in this region. These constraints are similar to those derived above, although some care should be taken in interpreting these numbers since the X-ray and radio flux densities are estimated over a rather large region ($1'' \sim 10 h_{50}^{-1} \text{ kpc}$). The magnetic field strength over such a large region (150 arcsec^2) may not be constant. That is, the radio surface brightness, spectral index, and minimum-energy magnetic field strength vary over this region, so it is almost certainly an oversimplification to use a total X-ray flux density, total radio flux density, single radio spectral index, and single minimum-energy magnetic field strength to describe this region. The constraints obtained are likely to be accurate for the extremities of the radio bridge region from which most of the radio emission originates and where most of the X-ray emission produced by IC scattering with CMB photons would originate.

The results obtained here, that the bridge magnetic field must be greater than about 0.3 of the minimum-energy value, are pushing against the upper bound of about 0.25 of the minimum-energy value required for consistency between the *Chandra* bounds on ambient gas densities and those predicted by WDW97 (discussed in § 5); given the roughness of the bounds, this trend would have to be confirmed with a larger sample of sources or with more *Chandra* data on these sources. It is interesting that the *Chandra* bounds on the ambient gas density tend to require a bridge magnetic field strength that is about a factor of 4 below the minimum-energy value, while the *Chandra* bounds on IC scattering from the bridge region tend to require a bridge magnetic field strength that is on the order of or greater than about 0.3 of the minimum-energy value. Since the upper and lower bounds are rather close, it will be necessary to consider the variations in bridge radio surface brightness, spectral index, and magnetic field strength when predicting bridge X-ray emission produced via IC scattering with CMB photons. Similarly, the X-ray data may be of sufficient quality to require a more detailed study of the radio-predicted ambient gas densities.

7. SUMMARY

We conclude that not all high-redshift 3CR radio sources are in rich cluster environments with luminous intracluster X-ray emission. In particular, a cluster like that which surrounds Cyg A is not present around 3C 254 or 3C 280. Even though the appearance of each radio source is consistent with the presence of a high-pressure environment, the luminosity and temperature of the surrounding gas could be significantly less than that of typical massive clusters of galaxies (more like that of a group or a poor cluster). These observations indicate that high-redshift radio sources are not necessarily the signposts of the presence of massive, virialized clusters of galaxies. Comparisons of predicted

ambient densities from radio observations to upper limits of the ambient density of a hot, confining medium suggest that $b = \frac{1}{4}$ models are yet consistent with the X-ray observations, while $b = 1$ models are ruled out under most assumptions. The sources 3C 254 and 3C 280 were selected from the WDW97 sample because of the confidence in their *ROSAT* properties; coincidentally, their radio-derived ambient densities were among the lowest in the WDW97 sample. An even more stringent test of the WDW97 predictions could be made with *Chandra* observations of radio sources with higher predicted densities.

The X-ray emission detected from the hot spot regions of 3C 280 and 3C 254 could be due to SSC emission, X-ray synchrotron emission, or IC scattering between relativistic electrons and external photon fields, such as the CMB. *HST*, radio, and X-ray data from the three hot spots detected by *Chandra* are consistent with X-ray synchrotron emission, although the radiative lifetimes of the relativistic electrons producing the X-ray emission are likely to be quite short and would require continuous acceleration in the hot spot region. SSC production of the X-ray emission can be very roughly estimated and would require magnetic field strengths of about 0.2–0.3 of the minimum-energy values; these extremely rough estimates could be significantly improved by employing high-resolution radio observations with broad frequency coverage. If SSC dominates, the hot spot *optical* emission detected in 3C 280 must have a separate origin from the X-ray emission unless the relativistic electron energy distribution extends to very low energies. Magnetic field strengths of 0.3–0.4 of the minimum-energy values suggest that IC scattering of CMB photons with relativistic electrons could also contribute to the total X-ray emission. As in the case of SSC, if IC scattering dominates the X-ray emission, the optical and X-ray hot spot emission must be produced by different processes unless the relativistic electron energy distribution extends to very low energies. We cannot conclude from these data whether one or more processes are responsible for the X-ray emission detected by *Chandra* and the possible optical emission detected by *HST* from the hot spots of these two radio sources.

The comparison of ambient gas densities predicted using radio data (WDW97) with the upper bounds provided by the *Chandra* data indicates that the magnetic field in the bridge regions of these sources must be below the minimum-energy values. The ambient gas densities predicted based on radio observations are consistent with the bounds placed using the *Chandra* data if the bridge magnetic field strength is about $\frac{1}{4}$ of the minimum-energy value. Meanwhile, *Chandra* constraints on IC-scattered X-rays from the bridge region tend to push the magnetic field strength in this region up relative to the minimum-energy value; the current bounds are marginally consistent with those obtained using the ambient gas density comparisons. Here, it is important to keep in mind that the radio bridge is a rather large region and may have a magnetic field strength that varies with bridge location. Longer *Chandra* exposures and/or *Chandra* data on similar sources that yield detections rather than bounds will be very helpful in determining the bridge magnetic field strengths and testing the reliability of the radio-determined ambient gas densities.

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REFERENCES

- Allen, S. W., et al. 2001, *MNRAS*, 324, 842
 Band, L., & Grindlay, J. E. 1985, *ApJ*, 298, 128
 Best, P. N. 2000, *MNRAS*, 317, 720
 Best, P. N., Longair, M. S., & Röttgering, H. J. A. 1997a, *MNRAS*, 286, 785
 ———. 1997b, *MNRAS*, 292, 758
 Borgani, S., et al. 2001, *ApJ*, 561, 13
 Bremer, M. H. 1997, *MNRAS*, 284, 126
 Brunetti, G., Bondi, M., Comastri, A., & Setti, G. 2002, *A&A*, 381, 795
 Burbidge, G. 1956, *ApJ*, 124, 416
 Carilli, C., Perley, P. A., Dreher, J. W., & Leahy, J. P. 1991, *ApJ*, 383, 554
 Carilli, C. L., Röttgering, H., van Ojik, R., Miley, G. K., & van Breugel, W. 1997, *ApJS*, 109, 1
 Crawford, C. S., Lehmann, I., Fabian, A. C., Bremer, M. N., & Hasinger, G. 1999, *MNRAS*, 308, 1159
 Daly, R. A. 1992a, *ApJ*, 386, L9
 ———. 1992b, *ApJ*, 399, 426
 ———. 2002, *NewA Rev.*, 46, 47
 De Vries, W. H., O'Dea, C. P., Barthel, P. D., Fanti, C., & Fanti, R. 2000, *AJ*, 120, 2300
 Dickey, J. M., & Lockman, F. J. 1990, *ARA&A*, 28, 215
 Donahue, M., & Voit, G. M. 1999, *ApJ*, 523, L137
 Donahue, M., Voit, G. M., Gioia, I., Luppino, G., Hughes, J. P., & Stocke, J. T. 1998, *ApJ*, 502, 550
 Donahue, M., Voit, G. M., Scharf, C. A., Gioia, I. M., Mullis, C. R., Hughes, J. P., & Stocke, J. T. 1999, *ApJ*, 527, 525
 Eke, V. R., Cole, S., Frenk, C. S., & Henry, J. P. 1998, *MNRAS*, 298, 1145
 Fabian, A. C., Crawford, C. S., Ettori, S., & Sanders, J. S. 2001, *MNRAS*, 322, L11
 Fixsen, D. J., Cheng, E. S., Gales, J. M., Mather, J. C., Shafer, R. A., & Wright, E. L. 1996, *ApJ*, 473, 576
 Guerra, E. J., Daly, R. A., & Wan, L. 2000, *ApJ*, 544, 659
 Hardcastle, M. J., Birkinshaw, M., & Worrall, D. M. 2001, *MNRAS*, 323, L17
 Hardcastle, M. J., & Worrall, D. M. 1999, *MNRAS*, 309, 969
 ———. 2000, *MNRAS*, 319, 562
 Harris, D. E., et al. 2000, *ApJ*, 530, L81
 Harvanek, M., & Stocke, J. T. 2002, *AJ*, 124, 1239
 Henry, J. P. 1997, *ApJ*, 489, L1
 ———. 2000, *ApJ*, 534, 565
 Leahy, J. P., & Williams, A. G. 1984, *MNRAS*, 210, 929
 Le Fèvre, O., Deltorn, J. M., Crampton, D., & Dickinson, M. 1996, *ApJ*, 471, L11
 Lehnert, M. D., et al. 1999, *ApJS*, 123, 351
 Liu, R., Pooley, G., & Riley, J. M. 1992, *MNRAS*, 257, 545
 Markevitch, M. 1998, *ApJ*, 503, 77
 McCarthy, P. J., van Breugel, W., Spinrad, H., & Djorgovsky, S. 1987, *ApJ*, 321, L29
 Mobasher, B. 2002, *HST Data Handbook*, Version 4.0 (Baltimore: STScI)
 Owen, F. N., & Puschell, J. J. 1984, *AJ*, 89, 932
 Pascarelle, S. M., Windhorst, R. A., Keel, W. C., & Odewahn, S. C. 1996, *Nature*, 383, 45
 Peebles, P. J. E., Daly, R. A., & Juskiewicz, R. 1989, *ApJ*, 347, 563
 Pentericci, L., Röttgering, H. J. A., Miley, G. K., McCarthy, P., Spinrad, H., van Breugel, W. J. M., & Macchetto, F. 1999, *A&A*, 341, 329
 Perley, R. A., & Taylor, G. B. 1991, *AJ*, 101, 1623
 Ridgway, S. E., & Stockton, A. 1997, *AJ*, 114, 511
 Sambruna, R. M., Eracleous, M., & Mushotzky, R. F. 1999, *ApJ*, 526, 60
 Sarazin, C. L. 1988, *X-Ray Emissions from Clusters of Galaxies* (Cambridge: Cambridge Univ. Press)
 Smith, D. A., Wilson, A. S., Arnaud, K. A., Terashima, Y., & Young, A. J. 2002, *ApJ*, 565, 195
 Spinrad, H., Djorgovski, S., Marr, J., & Aguilar, L. 1985, *PASP*, 97, 932
 Stark, A. A., Gammie, C. F., Wilson, R. W., Bally, J., Linke, R. A., Heiles, C., & Hurwitz, M. 1992, *ApJS*, 79, 77
 Wan, L., Daly, R. A., & Guerra, E. J. 2000, *ApJ*, 544, 671
 Wellman, G. F., Daly, R. A., & Wan, L. 1997, *ApJ*, 480, 79 (WDW97)
 Wilson, A. S., Young, A. J., & Shopbell, P. L. 2000, *ApJ*, 544, L27
 Worrall, D. M., Birkinshaw, M., Hardcastle, M. J., & Lawrence, C. R. 2001, *MNRAS*, 326, 1127
 Worrall, D. M., Lawrence, C. R., Pearson, T. J., & Readhead, A. C. S. 1994, *ApJ*, 420, L17